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The globular cluster system of NGC 6822

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ABSTRACT

We present a comprehensive analysis of the globular cluster (GC) system of the Local Group dwarf irregular galaxy NGC 6822. Our study is based on homogeneous optical and near-IR photometry, as well as long-slit spectroscopic observations which are used to determine new radial velocities for six GCs, two of which had no previous spectroscopic information. We construct optical-near-IR colour–colour diagrams and through comparison to simple stellar population models infer that the GCs have old ages consistent with being 9 Gyr or older, while their metallicities are in the range between $-1.6 \lesssim [\text{Fe}/\text{H}] \lesssim -0.4$. We conduct a kinematic analysis of the GC population and find tentative evidence for weak net rotation of the GC system, in the same sense as that exhibited by the underlying spheroid. The most likely amplitude of rotation is $\approx 10 \text{ km s}^{-1}$, approximately half the magnitude of the observed velocity dispersion. Finally, we use the GCs to estimate the dynamical mass of NGC 6822 within $\sim 11 \text{ kpc}$ and we formally find it to be in the range between $(3 \text{ and } 4) \times 10^9 M_{\odot}$. This implies an overall mass-to-light ratio in the range of $\sim 30\text{--}40$ and indicates that NGC 6822 is highly dark-matter-dominated. The mass and the corresponding mass-to-light ratio estimates are affected by various additional systematic effects due to limitations of the data and the model that are not necessary reflected in the formal uncertainties.

Key words: globular clusters: general – galaxies: individual: NGC 6822 – galaxies: kinematics and dynamics – Local Group.

1 INTRODUCTION

According to the Λ cold dark matter cosmological model, massive galaxies build up in part through the amalgamation and merger of smaller galaxies. The globular cluster (GC) systems of such massive galaxies, which are often quite rich, are also thought to be assembled, at least partially, from GCs donated by the smaller accreted galaxies (e.g. Côté, Marzke & West 1998; Côté et al. 2000). If this is indeed the case, GCs can act as valuable tracers of such events, while offering the potential to understand the formation history of their host galaxy. In their influential study, Searle &

Zinn (1978) uncovered important clues about the formation of the Milky Way through its GC system, namely that accretion played a significant part in the process. Modern studies of the Galactic GC system arrive at the same conclusion (e.g. Zinn 1993; Mackey & Gilmore 2004; Mackey & van den Bergh 2005; Marín-Franch et al. 2009; Dotter et al. 2010; Dotter, Sarajedini & Anderson 2011; Keller, Mackey & Da Costa 2012). Indeed the growth of our Galaxy and its GC system is still ongoing. Observations of the Sagittarius dwarf galaxy, which is currently being accreted on to the Milky Way while donating at least five GCs to the Galactic halo, provide the strongest evidence that this is the case (e.g. Bellazzini, Ferraro & Ibata 2003; Law & Majewski 2010). The work of our group has shown that a similar situation is also observed in M31. The halo of this galaxy features a number of prominent stellar streams

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with which many GCs are associated, bearing witness to an ample accretion history (e.g. Mackey et al. 2010b, 2013a, 2014; Huxor et al. 2011; Veljanoski et al. 2013b, 2014).

The study of dwarf irregular galaxies (dIrr) is a particularly important topic within the context of understanding galaxy formation. Being by far the most common galaxy type at high redshift (e.g. Stiavelli, Fall & Panagia 2004; van der Wel et al. 2011), dIrr systems could have provided a significant contribution towards the growth of massive galaxies at early times. In addition, in order to use GC as probes of galaxy assembly, it is crucial to understand the properties of GC systems in dwarf galaxies, and how they relate to the host galaxy properties and to the properties of GCs formed *in situ* around a massive galaxy. In this context, dIrr galaxies are especially helpful, since in the local Universe they are often found either in isolation or in groups without a massive dominating companion, making them excellent systems for studying the pristine properties of GC systems.

Lying in the Southern hemisphere with Galactocentric coordinates of $l = 25^\circ.4$, $b = -18^\circ.4$, NGC 6822 is one of the most intriguing dwarf galaxies in the Local Group. Discovered by Barnard (1884), this barred dIrr galaxy features a number of peculiar properties that has made it the focus of much attention over the years, even though its low Galactic latitude makes observations quite challenging. It has an absolute magnitude $M_V = -15.2$ and a half-light radius of ~ 0.5 kpc (Hodge 1977; Hodge et al. 1991). Located at a distance of 472 kpc (Górski, Pietrzyński & Gieren 2011), it is the closest dIrr galaxy save for the Magellanic Clouds. It does not appear to be associated with either the Milky Way or M31, and it has no other neighbouring companions. This galaxy contains a substantial spheroidal component, that has been traced via its constituent red giant branch stars out to a radius of at least ~ 5 kpc (Battinelli, Demers & Kunkel 2006). Demers, Battinelli & Kunkel (2006) used carbon stars in the spheroid as kinematic tracers, finding evidence for rotational support. In addition, NGC 6822 contains large quantities of H I gas forming a disc, the semimajor axis of which is positioned perpendicular to the stellar spheroid (e.g. de Blok & Walter 2000, 2006; Woldrake, de Blok & Walter 2003). The H I gas is also observed to be rotating, but at right angles to the rotation of the spheroid, showing that there are at least two different kinematic components present in NGC 6822. This is further supported by a recent study on planetary nebulae and H II regions by Flores-Durán et al. (2014), showing that the planetary nebulae have kinematics similar to the carbon stars, while the H II regions share the kinematics of the H I disc.

In recent years, the dedicated searches have substantially increased the number of GCs known around Local Group galaxies (e.g. Huxor et al. 2008; Veljanoski et al. 2013a), including NGC 6822. Hubble VII, originally discovered by Hubble (1925), was long considered the only old GC in NGC 6822. However, this changed when Hwang et al. (2011) discovered four new GCs in wide-field Canada–France–Hawaii Telescope (CFHT)/*MegaCam* imagery covering an area of 3×3 deg² around NGC 6822. Later, Huxor et al. (2013) re-examined this data and by complementing it with additional archival *Subaru* and CFHT imagery which expanded the observed footprint, uncovered three additional GCs, bringing the total number of members to 8. Fig. 1 shows the positions of the GCs superimposed on an image of NGC 6822, on which are also marked the contours of the H I distribution, ranging between 3.5 and 42×10^{20} cm⁻² (de Blok & Walter 2000), and the known extent of the red giant branch stars in the galaxy spheroid (Battinelli et al. 2006). It is notable that the GC system of NGC 6822 is quite spatially extended, with the most remote cluster, SC1, having a projected radius of nearly 11 kpc.

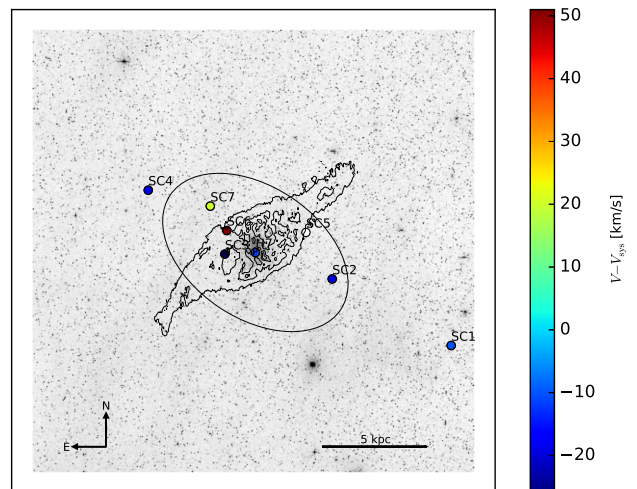


Figure 1. Positions of the eight known GCs overlaid on a Digital Sky Survey image of NGC 6822. The colours of the points correspond to their radial velocities, as measured in this work (see Section 5) and corrected for the systemic motion of NGC 6822. The ellipse shows the extent of the stellar spheroid traced by Battinelli et al. (2006), with a semimajor axis of 36 arcmin and an ellipticity of 0.36. Also shown are contours of the H I distribution map of de Blok & Walter (2000) ranging from 3.5 and 42×10^{20} cm⁻².

Furthermore, the GCs are arranged in a rather linear configuration, with an orientation aligned quite closely with the major axis of the spheroid (Hwang et al. 2011; Huxor et al. 2013). Five of the eight GCs in NGC 6822 are fairly extended in nature and thus resemble the extended clusters found in the haloes of M31 (e.g. Huxor et al. 2005, 2014) and M33 (Stonkutė et al. 2008). The coordinates, position angles, projected radii (R_{proj}), GC type and Galactic $E(B - V)$ extinction values (Schlegel, Finkbeiner & Davis 1998) of the clusters are listed in Table 1.

In this paper, we present a comprehensive analysis of the GC system in NGC 6822, based on uniform optical and near-IR photometry and long-slit spectroscopy. The photometry is used to constrain the ages and metallicities. Even though deeper optical data exist for some clusters (e.g. Hwang et al. 2011; Huxor et al. 2013), the uniformity of the data presented in this study ensures no systematic offsets are present in the age and metallicity estimates due to heterogeneous imaging. The spectroscopy yields new radial velocity measurements for six of the GCs, two of which had no previous velocity information. These data are also used to redetermine the dynamical mass and mass-to-light ratio of this galaxy.

2 THE DATA

2.1 Optical data

The optical imagery has been obtained as part of the *PanSTARRS* 3 π Survey (Kaiser et al. 2010; Chambers et al., in preparation). This survey has targeted three quarters of the sky ($\delta > -30^\circ$) in five optical bands, g_{P1} , r_{P1} , i_{P1} , z_{P1} , y_{P1} , with the 1.8 m *PanSTARRS* telescope located on Haleakala, Hawaii (Tonry et al. 2012). The optical system of *PanSTARRS* features a 1.4 Giga-pixel imager (Onaka et al. 2008; Tonry & Onaka 2009) with a field of view of 7 deg². The sky has been observed up to four times a year in each band, with individual exposures being between 30 and 45 s in length. The data are automatically processed in real time with the IMAGE PROCESSING PIPELINE (IPP; Magnier 2006). After the

Table 1. Coordinates, position angles, projected radii, GC type (compact or extended), the half-light radius (if known) and colour excess values due to Galactic interstellar reddening of the GCs in NGC 6822. The adopted centre of NGC 6822 is RA = 19^h44^m57^s.7, Dec. = −14°48′12″.

| ID | RA (J2000) (h m s) | Dec. (J2000) (d m s) | R_{proj} (kpc) | θ (°) | GC type | r_h (pc) | $E(B - V)$ (mag) | Reference |
|------------|-----------------------|-------------------------|----------------------------|-----------------|---------|----------------|---------------------|---------------------|
| Hubble-VII | 19 44 55.8 | −14 48 56.2 | 0.1 | 227 | C | 2.5 ± 0.1 | 0.24 | Hubble (1925) |
| SC1 | 19 40 11.9 | −15 21 46.6 | 10.7 | 244 | E | 14.0 ± 0.2 | 0.16 | Hwang et al. (2011) |
| SC2 | 19 43 04.5 | −14 58 21.4 | 4.1 | 250 | E | 11.5 ± 0.2 | 0.22 | Hwang et al. (2011) |
| SC3 | 19 45 40.2 | −14 49 25.8 | 1.4 | 94 | E | 7.5 ± 0.5 | 0.19 | Hwang et al. (2011) |
| SC4 | 19 47 30.4 | −14 26 49.3 | 6.0 | 60 | E | 13.8 ± 0.3 | 0.19 | Hwang et al. (2011) |
| SC5 | 19 43 42.3 | −14 41 59.7 | 2.7 | 290 | E | – | 0.22 | Huxor et al. (2013) |
| SC6 | 19 45 37.0 | −14 41 10.8 | 1.6 | 52 | C | – | 0.19 | Huxor et al. (2013) |
| SC7 | 19 46 00.7 | −14 32 35.0 | 3.0 | 43 | C | – | 0.21 | Huxor et al. (2013) |

standard reduction processes such as flat-fielding, the pipeline re-samples all images to a uniform pixel size of 0.25 arcsec and aligns them to the equatorial axes on to ‘skycells’ – regular patches on the sky of 6250 pixels across. The data are calibrated to each other self-consistently using partially overlapping exposures (Schlafly et al. 2012), yielding a calibration precision better than 10 mmag as measured against SDSS. For the work presented here, we use data from the processing version 2 (PV2) and 2×2 arcmin cutouts of stacked images of the GCs provided by the postage stamp server.

2.2 Near-IR data

The near-IR imaging of NGC 6822 was taken as part of a project to survey red stellar populations of Local Group galaxies. The data were obtained in 2008 October using the WFCAM instrument (Casali et al. 2007) mounted on United Kingdom Infrared Telescope (UKIRT). WFCAM has four individual detectors, each with a pixel scale of 0.4 arcsec, arranged such that four dithered pointings cover a square of 0.75 deg^2 when properly aligned. Our observations were conducted in three near-IR bands J , H and K . The data were processed with a pipeline created by the Cambridge Astronomy Survey Unit,¹ which carries out the standard dark current correction, flat-fielding, crosstalk removal, systematic noise and sky removal. The pipeline, described in Cioni et al. (2008) and Hodgkin et al. (2009), also performs full astrometric and photometric calibration, based on the 2MASS point source catalogue. The reduced images are stacked and microstepped to produce individual science frames. The finalized frames are not resampled, and thus they retain the original pixel scale of the detector. The nightly zero-point variation is < 1 per cent. Fig. 2 shows 60×60 arcsec wide cutouts of each GC hosted by NGC 6822 in the *PanSTARRS1* g_{PI} and UKIRT/WFCAM K bands. Note that SC1 falls outside the near-IR coverage and thus no thumbnail is shown, while SC5 is only marginally detected in both optical and near-IR data sets.

2.3 Spectroscopic data

Moderate resolution spectra for six GCs were acquired in the summer of 2009 using the ISIS instrument mounted on the 4.2-m William Herschel Telescope (WHT), the RC-spectrograph installed on the 4-m Kitt-Peak National Observatory (KPNO) telescope, and the Deep Extragalactic Imaging Multi-Object Spectrograph (DEIMOS) spectrograph used by the 10-m W. M. Keck telescope. Several exposures

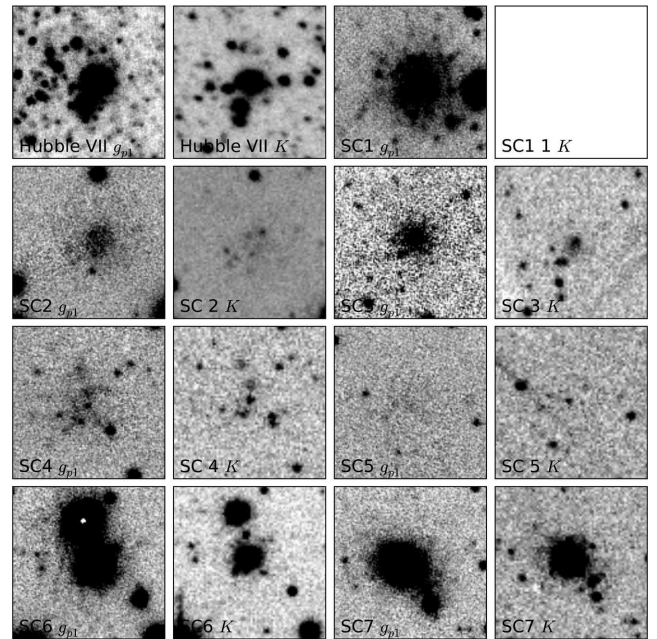


Figure 2. g_{PI} and K -band images of the GCs in NGC 6822. Each image is 60×60 arcsec wide. North is up and east is left. SC1 falls outside the near-IR survey and hence no K -band image is shown, while SC5 is only marginally detected in both passbands.

of each GC were obtained, with individual integration times determined according to the brightness of each target. Our observing log is shown in Table 2.

2.3.1 WHT and KPNO data

The ISIS instrument comprises two individual detectors, attached to separate ‘arms’ of the spectrograph, that independently sample different wavelength ranges. The blue arm covers a wavelength range between ~ 3500 and 5100 Å with a dispersion of 0.4 Å pixel^{-1} , selected via the usage of the EE12 camera along with the R600B grating. The resolving power is $R \sim 1500$. Conversely, the red arm uses the RED-PLUS detector together with the R600R grating to select the wavelength range between ~ 7400 and 9200 Å with a dispersion of $0.49 \text{ Å pixel}^{-1}$. The resolving power in this case is $R \sim 2700$. A similar set up was used for the KPNO observations. The T2KB detector, in conjunction with the KPC007 grating were employed in order to select the wavelength range between ~ 3500 and 6500 Å with a dispersion of $0.139 \text{ Å pixel}^{-1}$ and a resolving

¹ <http://casu.ast.cam.ac.uk>

Table 2. Log of spectroscopic observations for GCs around NGC 6822.

| ID | Date of obs. | Number of exposures | Integration time per exposure (s) | Telescope | Principal investigator |
|------------|---------------|---------------------|-----------------------------------|-------------|------------------------|
| Hubble VII | 14/08/2009 | 3 | 900 | <i>KPNO</i> | A. Ferguson |
| SC1 | 16/08/2009 | 3 | 1800 | <i>KPNO</i> | A. Ferguson |
| SC3 | 17/08/2009 | 3 | 1800 | <i>KPNO</i> | A. Ferguson |
| SC6 | 15/08/2009 | 3 | 900 | <i>KPNO</i> | A. Ferguson |
| SC7 | 15/08/2009 | 3 | 600 | <i>KPNO</i> | A. Ferguson |
| SC2 | 16–17/08/2009 | 3 | 1800 | WHT | A. Huxor |
| SC7 | 16/08/2009 | 2 | 600 | WHT | A. Huxor |
| SC1 | 26/06/2009 | 3 | 300 | Keck | J. Hurley |
| SC2 | 26/06/2009 | 3 | 300 | Keck | J. Hurley |
| SC6 | 26/06/2009 | 3 | 300 | Keck | J. Hurley |
| SC7 | 26/06/2009 | 3 | 420 | Keck | J. Hurley |

power of $R \sim 1300$. The spectra obtained with both telescopes were not binned in either the spatial or the wavelength directions. The typical S/N of the spectra is $\sim 4\text{--}20$ per Å.

The reduction of the spectroscopic data was carried out in IRAF.² The basic reduction steps – bias and overscan subtraction, flat-fielding, illumination correction – were done with appropriate tasks in the CCD package. Having completed the initial reduction steps on the two-dimensional frames, the *apall* task in the *KPNOSLIT* package was used to extract one-dimensional spectra from them. We set the extraction aperture to be between 2 and 4 arcsec. With the same task, we interactively selected appropriate background sky regions, which we fit with a second-order Chebyshev polynomial and subtracted. The spectra were then traced using a third-order cubic spline function, and extracted using the optimal variance weighting option in *apall*.

We established wavelength calibration of the one-dimensional frames via Cu-Ne-Ar and He-Ne-AR lamps for WHT and *KPNO* spectra, respectively. Comparison ‘arcs’ were acquired before and after each cluster exposure. We extracted the arc spectra with an identical strategy as the target GCs they were used to calibrate. Using the *identify* task, we pinpointed ~ 90 *ISIS* blue, ~ 25 *ISIS* red and ~ 50 RC lines in the arc spectra and fitted a dispersion solution with a third-order cubic spline function. The root-mean-square residuals of the fits are 0.05 ± 0.01 , 0.02 ± 0.01 and 0.08 ± 0.01 Å for the data obtained with the *ISIS* blue, *ISIS* red and RC instruments, respectively. Since two arc spectra were observed for each target exposure, the resulting wavelength solutions were averaged before being applied to the appropriate GC via the *dispcor* task. To assess the reliability of the wavelength calibration, we measured the positions of sky emission lines in the sky spectra, which are a by-product of the *apall* extraction. We found that the wavelength calibration is accurate to 0.08 Å for the WHT and 0.1 Å for the *KPNO* data.

Given that all but one cluster were observed as a series of separate exposures with each instrument, these were stacked using an inverse variance weighting technique (e.g. Veljanoski et al. 2014). Finally, we continuum subtracted the spectra for the purpose of measuring radial velocities. Examples of the fully reduced spectra obtained with each instrumental set up are shown on Fig. 3. The displayed spectra are continuum normalized rather than continuum subtracted in order to preserve the relative strengths of the absorption lines for better visualization.

2.3.2 Keck data

We used the DEIMOS in longslit mode to observe four NGC 6822 GCs (Swinburne Keck programme 2009A_W006D, PI Hurley). The longslit mode of DEIMOS is achieved by inserting a standard single slit mask into the optical path of the instrument. For all observations with this spectrograph, we used a mask featuring a slit 0.8 arcsec wide. The DEIMOS camera comprises of eight CCD chips arranged in a 4×2 mosaic pattern, with the long axis of this array oriented in the spatial direction and the short axis in the dispersion direction. This setup essentially produces two spectra simultaneously for a given compact target – one on a ‘blue’ CCD and one on a ‘red’ CCD. We selected a central wavelength of 6040 Å via the 1200 l mm^{-1} grating, resulting in a dispersion of 0.33 Å pixel^{-1} . The resolving power of the spectra observed with the red CCD is $R \sim 4700$.

To reduce observations of a given GC, we first identified the relevant CCD chips covered by the blue and red spectra. We isolated these and reduced them separately; the remaining six CCD chips were discarded, as they contained no signal coming from our target. Following this, the standard reduction steps such as bias and overscan subtraction, and flat-fielding, were undertaken with appropriate routines in the CCD package in IRAF. Inspecting the reduced two-dimensional frames, significant geometrical distortions in the dispersion of the target objects are clearly evident, as well as distortions in the sky lines. To rectify these, we simply wavelength calibrated the two-dimensional frames prior to extracting a one-dimensional spectrum.

The wavelength solution is based on a Ne-Ar-Kr-Xe lamp, taken during the baseline calibrations of the instrument. For the frames containing the redder part of the spectrum, we identified 23 lines in the arc spectrum via the *identify* and *reidentify* tasks. The two-dimensional wavelength solution was then determined with the *fitcoords* task located in the *TWODSPEC* package. The root-mean-square residuals of the fit were 0.02 ± 0.01 Å. We next applied the wavelength solution to the appropriate frames via the *transform* routine found in the same package. Inspecting the wavelength calibrated frames, one sees that the distortions in the spatial axis have disappeared, and the skylines are perfectly straight in a direction perpendicular to the dispersion axis. This is a good indicator that the wavelength calibration is reliable. For clarity, the wavelength range of our spectra observed with the red chip spanned $\sim 6000\text{--}7250$ Å.

Unfortunately, we were unable to wavelength calibrate the frames containing the blue part of the spectra. This is because the arc spectrum observed at these wavelengths contained only three reliable lines, all of which sit extremely close to the edge of the frame.

² IRAF is distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

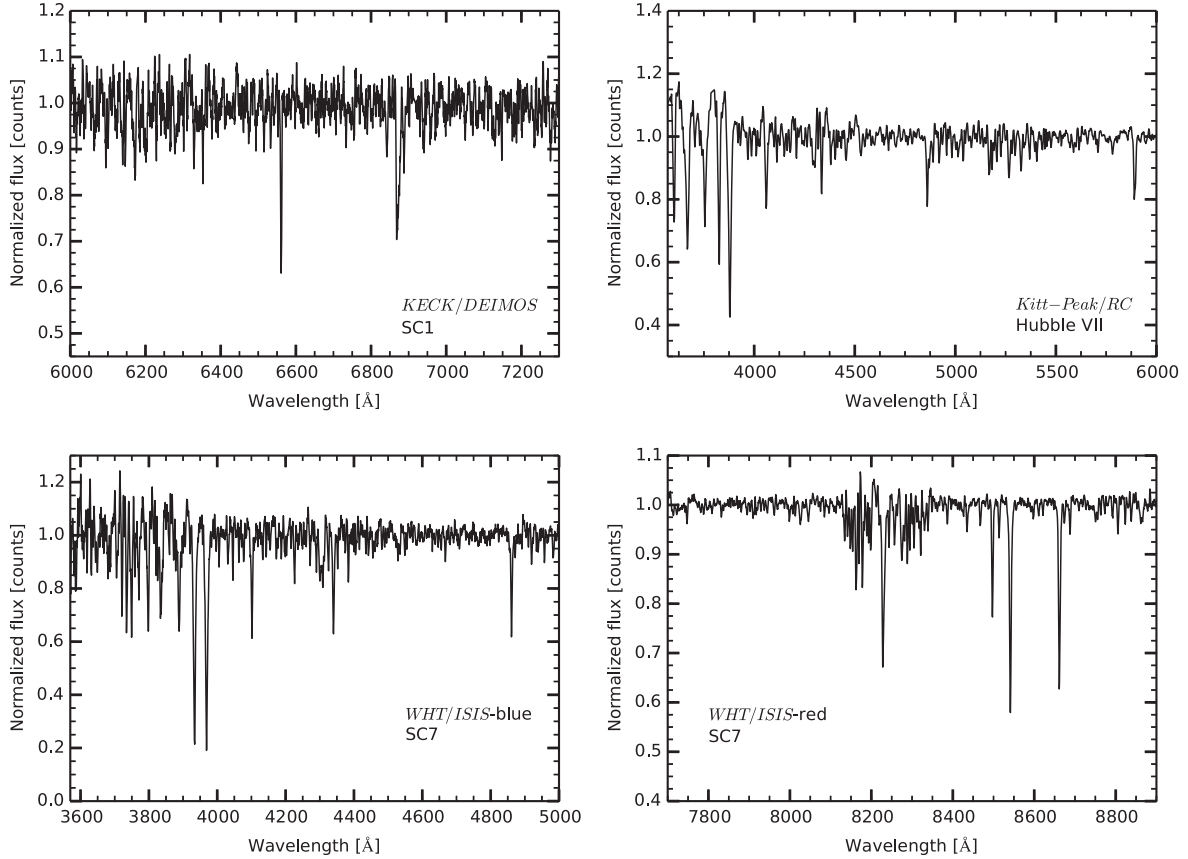


Figure 3. Examples of fully reduced, continuum-normalized spectra of GCs in NGC 6822.

This is a well-known calibration issue for data taken in this specific wavelength range (albeit one that is nowadays possible to circumvent). Even though this problem forced us to proceed further using only the ‘red’ frames, it did not hinder any of the velocity estimation and kinematic analysis which we present in the later sections of this paper.

From the appropriate two-dimensional frames, we extracted one-dimensional spectra and stacked these in the same manner as for the 4-m data, with the only difference being the width of the extraction aperture, which for the Keck data we set to 1.1 arcsec. As before, we used the sky spectra associated with the extracted one-dimensional spectra to assess the accuracy of the wavelength calibration. We find that the wavelength calibration is typically accurate to 0.06 Å. An example of a fully reduced, wavelength calibrated Keck spectrum is shown in Fig. 3.

3 PHOTOMETRY

We determined the total magnitudes of each GC in NGC 6822 via aperture photometry in all eight bands, g_{P1} , r_{P1} , i_{P1} , z_{P1} , y_{P1} , J , H , and K . The photometric measurements were conducted using the *phot* task in IRAF. Accurate centres for the classical, i.e. compact, GCs were established with the centroid algorithm within *phot*. Since the astrometric calibration is identical for the entire *PanSTARRS1* data set, we performed the centring in the r_{P1} band imaging, since the clusters are best seen there. For the diffuse, extended clusters, we determined the centre via visual inspection. Offsets from a nearby

bright star were used to accurately redetermine the centre of each cluster through the rest of the data.

We used circular apertures to sum the total flux coming from each cluster. The aperture size for a given cluster was chosen such that the aperture enclosed the entire extent of the target; these sizes are listed in Table 3. The background sky was determined between two concentric circular annuli surrounding the photometric aperture. We carefully placed these in a way so as to exclude unwanted contaminants such as extended background galaxies or foreground Milky Way stars. The reported magnitude uncertainties, which also include the variation in the sky, are formally determined by the *phot* task.

All our photometric measurements were zero-point calibrated and corrected for atmospheric extinction, and then subsequently corrected for foreground reddening using $E(B - V)$ values from Schlegel et al. (1998). We calculated the extinction coefficients for the *PanSTARRS1* bands using the relations presented in Tonry et al. (2012):

$$Ag_{P1}/E(B-V) = 3.613 - 0.0972(g_{P1} - i_{P1}) + 0.0100(g_{P1} - i_{P1})^2 \quad (1)$$

$$Ar_{P1}/E(B-V) = 2.585 - 0.0315(g_{P1} - i_{P1}) \quad (2)$$

$$Ai_{P1}/E(B-V) = 1.908 - 0.0152(g_{P1} - i_{P1}) \quad (3)$$

$$Az_{P1}/E(B-V) = 1.499 - 0.0023(g_{P1} - i_{P1}) \quad (4)$$

$$Ay_{P1}/E(B-V) = 1.251 - 0.0027(g_{P1} - i_{P1}). \quad (5)$$

Table 3. Optical and near-IR photometry of all GCs hosted by NGC 6822. The superscript (1) denotes that all magnitudes for this cluster are lower limits (i.e. upper limits to the luminosity).

| ID | Aperture (arcsec) | g_0 (mag) | r_0 (mag) | i_0 (mag) | z_0 (mag) | y_0 (mag) | J_0 (mag) | H_0 (mag) | K_0 (mag) |
|------------------|----------------------|------------------|------------------|------------------|------------------|------------------|------------------|------------------|------------------|
| Hubble VII | 3.7 | 15.46 ± 0.02 | 14.99 ± 0.02 | 14.72 ± 0.02 | 14.58 ± 0.03 | 14.47 ± 0.04 | 13.44 ± 0.01 | 12.98 ± 0.02 | 12.86 ± 0.01 |
| SC1 | 7.0 | 16.62 ± 0.02 | 16.19 ± 0.02 | 15.97 ± 0.02 | 15.83 ± 0.03 | 15.72 ± 0.04 | — | — | — |
| SC2 | 8.0 | 17.64 ± 0.02 | 16.98 ± 0.02 | 16.80 ± 0.02 | 16.63 ± 0.03 | 16.57 ± 0.05 | 15.53 ± 0.03 | 15.11 ± 0.03 | 14.86 ± 0.03 |
| SC3 | 3.0 | 18.79 ± 0.02 | 18.35 ± 0.02 | 18.08 ± 0.02 | 17.98 ± 0.03 | 17.77 ± 0.04 | 16.93 ± 0.04 | 16.44 ± 0.04 | 16.32 ± 0.05 |
| SC4 | 8.0 | 17.94 ± 0.02 | 17.37 ± 0.02 | 17.00 ± 0.02 | 16.84 ± 0.03 | 16.78 ± 0.02 | 15.74 ± 0.04 | 15.31 ± 0.04 | 15.11 ± 0.04 |
| SC5 ¹ | 7.5 | ≥ 19.3 | ≥ 18.6 | ≥ 18.2 | ≥ 18.2 | ≥ 17.8 | ≥ 18.0 | ≥ 16.6 | ≥ 16.4 |
| SC6 | 4.7 | 15.76 ± 0.02 | 15.26 ± 0.02 | 15.03 ± 0.01 | 14.91 ± 0.03 | 14.83 ± 0.02 | 13.89 ± 0.01 | 13.42 ± 0.02 | 13.30 ± 0.01 |
| SC7 | 5.0 | 15.24 ± 0.02 | 14.60 ± 0.02 | 14.29 ± 0.01 | 14.12 ± 0.03 | 13.99 ± 0.02 | 12.93 ± 0.01 | 12.39 ± 0.02 | 12.30 ± 0.01 |

Table 4. Photometry of the NGC 6822 GCs converted to the Johnson/Cousins system using equations 6a and 6b taken from Tonry et al. (2012). For convenience, the $(V - I)_0$ colours are also listed. The superscript is as in Table 3.

| ID | V_0 (mag) | I_0 (mag) | $(V - I)_0$ (mag) | $M_{V,0}$ (mag) |
|------------------|------------------|------------------|----------------------|--------------------|
| Hubble VII | 15.05 ± 0.02 | 14.18 ± 0.03 | 0.87 ± 0.04 | -8.3 |
| SC1 | 16.30 ± 0.02 | 15.48 ± 0.03 | 0.82 ± 0.04 | -7.1 |
| SC2 | 17.15 ± 0.03 | 16.25 ± 0.03 | 0.90 ± 0.04 | -6.2 |
| SC3 | 18.44 ± 0.04 | 17.57 ± 0.03 | 0.87 ± 0.05 | -4.9 |
| SC4 | 17.52 ± 0.03 | 16.47 ± 0.02 | 1.05 ± 0.04 | -5.9 |
| SC5 ¹ | ≥ 18.8 | ≥ 17.6 | ~ 0.76 | ≥ -4.6 |
| SC6 | 15.38 ± 0.02 | 14.51 ± 0.02 | 0.87 ± 0.03 | -8.0 |
| SC7 | 14.77 ± 0.02 | 13.75 ± 0.02 | 1.02 ± 0.03 | -8.6 |

Schlafly & Finkbeiner (2011) recalibrated the $E(B - V)$ from Schlegel et al. (1998) and recommend that the *PanSTARRS1* reddening coefficients be further multiplied by a factor of 0.88, which we included in the magnitude derivation. Finally, Table 3 lists the extinction-corrected total magnitudes of all GCs in NGC 6822.

It is important to note that SC5 is at best marginally detected in the imaging, and thus the measured values represent lower limits to its magnitudes in all bands (i.e. upper limits to its luminosity). The extended cluster SC1 is the most remote object in this system and is located outside the coverage of the near-IR data.

To allow for comparison with other studies, we convert the photometry measurements from the *PanSTARRS1* system to the classical Johnson/Cousins system using equations derived in Tonry et al. (2012),

$$V - r_{P1} = 0.005 + 0.462(g - r)_{P1} + 0.013(g - r)_{P1}^2 \quad (6a)$$

$$I - i_{P1} = -0.366 - 0.136(g - r)_{P1} - 0.018(g - r)_{P1}^2. \quad (6b)$$

The conversion to the Johnson V and Cousins I bands adds an additional 0.012 and 0.017 mag to the uncertainty, respectively. The results of the conversions are listed in Table 4. This table also lists the absolute magnitudes $M_{V,0}$, calculated assuming each cluster has the same distance of 472 kpc as the centre of NGC 6822. In Table 5, we compare our photometric measurements to those in previously published studies. In their discovery paper, Huxor et al. (2013) presented optical photometric for SC6 and SC7, measured on CFHT/MegaCam archival images. There is an excellent agreement between the photometry presented in this work to that in Huxor et al. (2013), with the measurements being less than one standard deviation apart. Comparing the optical $(V - I)_0$ colour from our

Table 5. $(V - I)_0$ colour comparison between the present study to those of Huxor et al. (2013) and Hwang et al. (2011).

| ID | $(V - I)_0$ This work | $(V - I)_0$ Huxor et al. (2013) | $(V - I)_0$ Hwang et al. (2011) |
|------------|--------------------------|------------------------------------|------------------------------------|
| Hubble-VII | 0.87 ± 0.04 | — | — |
| SC1 | 0.82 ± 0.04 | — | 0.85 ± 0.03 |
| SC2 | 0.90 ± 0.04 | — | 0.94 ± 0.03 |
| SC3 | 0.87 ± 0.05 | — | 1.31 ± 0.03 |
| SC4 | 1.05 ± 0.04 | — | 1.12 ± 0.03 |
| SC5 | ~ 0.76 | — | — |
| SC6 | 0.87 ± 0.03 | 0.84 ± 0.03 | — |
| SC7 | 1.02 ± 0.03 | 1.05 ± 0.03 | — |

study to those published in Hwang et al. (2011), one finds mixed results. While there is an excellent agreement between the colour measurements of our present study and that of Hwang et al. (2011) for SC1 and SC2, there is a poor agreement regarding the colour values of SC4, and even a strong disagreement in the case of SC3. In addition, the colours presented by Hwang et al. (2011) are always redder than those in this study.

4 AGES AND METALLICITIES

The best way of determining metallicities and ages of GCs is through the analysis of deep colour-magnitude diagrams of their resolved stellar populations. Second to that, high-quality spectroscopic data can also be used to derive the metal content and age of GCs. However, not all of our spectra have sufficient high S/N to reliably determine these quantities. Broad-band colour-colour diagrams using a combination of optical-optical and optical-near-IR colours are another way to estimate ages and metallicities of GCs, and require only the existence of homogeneous well-calibrated photometry (e.g. Puzia et al. 2002; Hempel et al. 2003; Veljanoski et al. 2013a; Muñoz et al. 2014). The mechanism behind this is that both $(g - i)_0$ and $(g - K)_0$ have similar sensitivity to the stars located near the main-sequence turn off point, which are most sensitive to age, but $(g - K)_0$ is considerably more affected by the temperature of the red giant branch which is more dependent on metallicity than age. Hence, comparing where GCs on such a plot sit relative to simple stellar population model tracks allows a crude estimate of age and metal content of each cluster.

In Fig. 4, we show a $(g_{P1} - i_{P1})_0$ versus $(g_{P1} - K)_0$ colour-colour diagram constructed with the photometry presented in Table 3. Integrated magnitudes, in appropriate filters, of simple stellar populations having several discrete ages are also overlaid

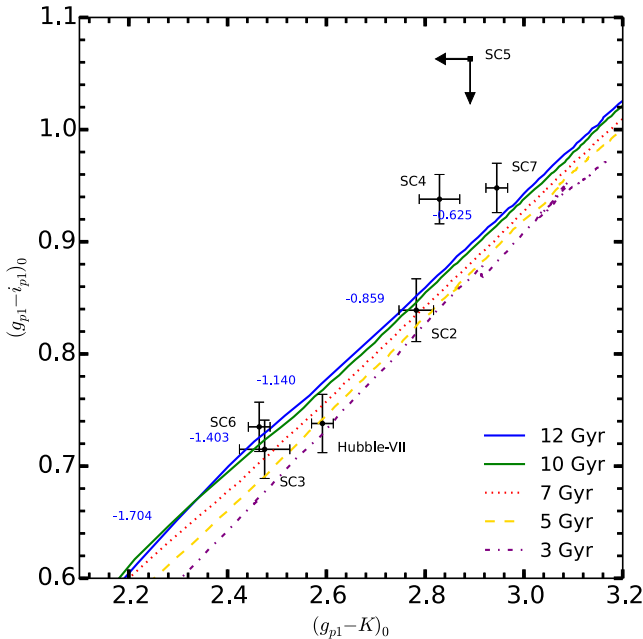


Figure 4. $(g_{p1} - i_{p1})_0$ versus $(g_{p1} - K)_0$ colour-colour diagram of the GCs in NGC 6822, overlaid on top of the simple stellar population isochrones derived from the PARSEC1.1 models (Bressan et al. 2012). Also shown are $[\text{Fe}/\text{H}]$ values along the 12 Gyr isochrone. SC5 is represented by a square since this cluster is a marginal detection in the imaging. The clusters with reliable photometry are all found to have old ages.

on the same plot. They are derived from the PARSEC1.1 models (Bressan et al. 2012), and span the metallicity range of $-2.2 < [\text{Fe}/\text{H}] < +0.6$ dex. Formally, the diagram shows that the clusters have old ages spanning the range between ~ 5 and 12 Gyr, but they are all consistent with being 9 Gyr or older. This is in agreement with the spectroscopic study by Hwang et al. (2014). The only exception is the cluster SC5, which being only marginally detected at best in both the optical and near-IR imaging, does not lie near the model tracks and thus its age and metallicity cannot be reliably constrained. For completeness, it is shown as a filled square in Fig. 4.

We estimate the metallicities of the clusters via two different methods. In both methods, we ignore SC5 since this cluster does not have reliable photometry. In the first method, the metallicity of each cluster is determined directly from Fig. 4. To do this, we

projected the position of each GC on to a grid of 12 Gyr isochrones in colour-colour space and determined which track lay closest to the GC. The 12 Gyr isochrone set is well sampled in metallicity and thus no interpolation is needed for extracting the correct metallicities. The upper and lower uncertainties are determined in the exactly the same way, considering the bluest and reddest possible combination of colours of each clusters, respectively, as allowed by the colour uncertainties. The metallicities of the clusters determined in this way are presented in the first column in Table 6.

A different method of estimating GC metallicities is through empirical colour-metallicity relations. Here, we use the relation derived by Kissler-Patig, Brodie & Minniti (2002), calibrated using 129 Milky Way and M31 GCs that have $E(B - V) < 0.27$, to set additional constraints on the metallicities of the NGC 6822 GCs. The relation is valid over a metallicity range of $-2.3 < [\text{Fe}/\text{H}] < -0.2$ dex, and it has the form

$$[\text{Fe}/\text{H}] = (V - K)_0 \times 1.82 \pm 0.11 - 5.52 \pm 0.26 \quad (7)$$

with an rms of 0.29 dex. Note that such a linear relation between colour and metallicity is a crude assumption (Cantiello & Blakeslee 2007).

The metallicities obtained via equation (7) are also shown in Table 6. It is apparent that there is a modest discrepancy between the estimated metallicities via the two different methods. Although all measurements, except that for SC2, formally agree with each other, a clear and nearly constant offset of ~ 0.37 dex is present. This mismatch is most likely due to systematics in the stellar evolutionary models. Given how rapidly these models progress, at the moment we find the metallicities estimated via the empirical relation more reliable.

Table 6 also lists the $[\text{Fe}/\text{H}]$ estimates from the spectroscopic study of Hwang et al. (2014), as well as spectroscopic metallicities for Hubble VII from earlier studies (Cohen & Blakeslee 1998; Chandar, Bianchi & Ford 2000). In general, the photometric metallicities derived here are considerably higher than those presented by Hwang et al. (2014), with only SC3 being in formal agreement. However, no uniform offset is present between these two sets of measurements. Furthermore, while our photometric $[\text{Fe}/\text{H}]$ value for Hubble VII is also higher than those derived in Cohen & Blakeslee (1998) and Chandar et al. (2000), the measurements agree within the reported uncertainties.

The reason for the discrepancy between the results presented in this contribution and those by Hwang et al. (2014) is unclear. It is known that NGC 6822 contains young and more metal-rich

Table 6. $[\text{Fe}/\text{H}]$ values for the GCs with reliable photometry in NGC 6822. Columns refer to: (a) values obtained directly from Fig. 4 by determining the projection of each cluster on to the 12 Gyr isochrone; (b) values obtained via equation (7), an empirical colour-metallicity relation calibrated on the M31 GC system (Kissler-Patig et al. 2002); spectroscopic metallicity from (c) Hwang et al. (2014); (d) Cohen & Blakeslee (1998); (e) Chandar et al. (2000).

| ID | $[\text{Fe}/\text{H}](a)$ (dex) | $[\text{Fe}/\text{H}](b)$ (dex) | $[\text{Fe}/\text{H}](c)$ (dex) | $[\text{Fe}/\text{H}](d)$ (dex) | $[\text{Fe}/\text{H}](e)$ (dex) |
|------------|------------------------------------|------------------------------------|------------------------------------|------------------------------------|------------------------------------|
| Hubble-VII | $-1.14^{+0.04}_{-0.09}$ | -1.5 ± 0.4 | -2.34 ± 0.03 | -1.95 ± 0.15 | -2.0 ± 0.25 |
| SC1 | — | — | -2.00 ± 0.04 | — | — |
| SC2 | $-0.84^{+0.04}_{-0.06}$ | -1.3 ± 0.4 | -2.53 ± 0.06 | — | — |
| SC3 | $-1.34^{+0.06}_{-0.15}$ | -1.7 ± 0.4 | -1.52 ± 0.06 | — | — |
| SC4 | $-0.75^{+0.06}_{-0.05}$ | -1.1 ± 0.4 | -2.53 ± 0.08 | — | — |
| SC5 | — | — | — | — | — |
| SC6 | $-1.40^{+0.07}_{-0.04}$ | -1.7 ± 0.3 | — | — | — |
| SC7 | $-0.61^{+0.03}_{-0.04}$ | -1.0 ± 0.4 | — | — | — |

Table 7. Heliocentric radial velocities for the GCs around NGC 6822 as observed from each instrument. Also shown are the combined radial velocities which are then used for the kinematic analysis. Note that SC7 was used as a velocity template for the Keck data set. For reference, the systemic velocity of NGC 6822, as well as velocity information available in the literature are also shown.

| ID | <i>KPNO</i> v (km s ⁻¹) | WHT _{blue} v (km s ⁻¹) | WHT _{red} v (km s ⁻¹) | <i>KECK</i> v (km s ⁻¹) | combined v (km s ⁻¹) | Literature v (km s ⁻¹) | Reference for literature data |
|------------|--|--|---|--|---------------------------------------|---|---|
| NGC 6822 | – | – | – | – | – | –57 | McConnachie (2012) |
| Hubble VII | -68 ± 12 | – | – | – | -68 ± 12 | $-52; -65 \pm 20$ | Cohen & Blakeslee (1998); Hwang et al. (2014) |
| SC1 | -69 ± 16 | – | – | -67 ± 4 | -67 ± 4 | -61 ± 20 | Hwang et al. (2014) |
| SC2 | – | -72 ± 12 | -80 ± 7 | -74 ± 5 | -76 ± 4 | -106 ± 31 | Hwang et al. (2014) |
| SC3 | -83 ± 14 | – | – | – | -83 ± 14 | -71 ± 17 | Hwang et al. (2014) |
| SC4 | – | – | – | – | – | $-115 \pm 58; -75 \pm 3$ | Hwang et al. (2014); Mackey et al. (in preparation) |
| SC5 | – | – | – | – | – | – | |
| SC6 | -16 ± 13 | – | – | -5 ± 3 | -6 ± 3 | – | |
| SC7 | -42 ± 17 | -36 ± 11 | -37 ± 2 | template | -37 ± 2 | – | |

populations of stars, and they could be the cause of the observed results if they were contaminating the photometric apertures. However, repeating the photometry with smaller apertures in order to exclude possible contaminants did not change the results in Fig. 4. In addition, the ages are recovered to be old which argues against younger stellar populations contaminating the derived total magnitudes. As already noted, the optical ($V-I$) colours presented by Hwang et al. (2011) are always redder than those in this work. It is clear from examining Fig. 4 that neither our nor their optical colours are consistent with their low spectroscopic metallicities, at least when interpreted using the PARSEC1.1 models. A colour–magnitude diagram analysis of the resolved stellar populations in these GCs would certainly help to resolve this issue, and might even indicate the reason for the mismatch.

Another effect to be considered is the internal reddening within NGC 6822 itself. As discussed earlier, this galaxy contains a significant amount of gas and dust, which, if not properly accounted for could cause the clusters to appear redder and thus more metal-rich than they actually are. Looking at Fig. 1 this seems to be unlikely since many GCs are located outside the extent of the H I disc, including several which are inferred to have high metallicities.

5 KINEMATICS

5.1 Radial velocity measurements

We derived heliocentric radial velocities from the spectra obtained with the 4-m class telescopes via a customized χ^2 minimization routine, using a number of appropriate radial velocity standard stars and bright M31 GCs as templates (see Veljanoski et al. 2014). The derived velocity from each science spectrum is the average of all minimizations between that spectrum and the applied velocity templates. A full description of the minimization process, as well as the radial velocity standard stars and clusters is presented in Veljanoski et al. (2014). The χ^2 minimization method is similar to the more frequently used cross-correlation technique (Tonry & Davis 1979), with the advantage that it uses the uncertainties in the spectra of both the target and template objects, which helps to weed out spurious peaks in the χ^2 functions and reduces the resulting velocity uncertainties.

During the course of the Keck observations no radial velocity standard stars were observed. However, the bright cluster SC7 was observed with both the *KPNO* and the WHT telescopes, essentially providing three independent velocity estimates that all agree with each other very well (one of these comes from *KPNO*, while two come from the red and blue arms of the *ISIS* spectrograph on the

WHT). Thus, we use our Keck observation of this cluster as a radial velocity template against which the other Keck spectra were matched, and adopt as its velocity the error weighted mean of the three independent measurements obtained from the 4-m data sets. The same χ^2 algorithm was used to estimate the radial velocity of the Keck targets.

From all these data, we have up to three independent velocity estimates for six of the eight GCs in NGC 6822, two of which (SC6 and SC7) possess no previous velocity information. Table 7 lists the velocity estimates coming from the different data sets and their associated uncertainties. Since these all agree within the measurement uncertainties, we combine them into a single value by performing an error weighted average and use these values for the forthcoming analysis. For comparison, Table 7 also lists the velocity information available in the literature, as well as the systemic velocity of NGC 6822. Our measurements are generally in good agreement with those from the literature, while possessing considerably smaller uncertainties.

Of the remaining two clusters, SC5 is sufficiently faint that to date there have been no successful velocity measurements made for this object. Hwang et al. (2014) present a velocity measurement for SC4 (see Table 7); however the uncertainties on their estimate are large. Although we did not obtain a longslit spectrum of SC4 during any of our observing runs, individual stars within this object were targeted with *DEIMOS* in multislit mode during our Keck run. The purpose of these observations was to investigate the internal kinematics of this rather extended cluster, and this analysis will be presented elsewhere (Mackey et al., in preparation). Here, we adopt the outcome of a computation to determine the most-likely radial velocity of this cluster given individual measurements for seven member stars, $v = -75 \pm 3$ km s⁻¹.

5.2 Bayesian kinematic modelling

Constraining the kinematic properties of the GCs around NGC 6822 is one of the main goals of this paper. Unlike in massive galaxies, the GC systems of dwarfs are typically populated by only a few members, so it is important to use a method which will extract the maximum information from the sparse data that are available. Following the example from Veljanoski et al. (2014), we work in the Bayesian framework, which enables us to construct a kinematic model and deduce the most likely values for the quantities of interest.

Given the known rotation of both the H I disc and the spheroid (as traced with carbon stars), we particularly want to test whether the GC system of NGC 6822 exhibits any measurable degree of

rotation. We thus construct a kinematic model \mathcal{M} which features both rotational and velocity dispersion components. The rotation component is constructed following the prescription from Côté et al. (2001),

$$v_p(\theta) = v_{\text{sys}} + A \sin(\theta - \theta_0), \quad (8)$$

where v_p are the measured radial velocities, θ and θ_0 are the position angles of the GCs and the GC rotation axis, respectively, and A is the rotation amplitude. The systemic radial motion of the GC system taken to be the heliocentric motion of NGC 6822, is labelled as v_{sys} . As elaborated in Côté et al. (2001), this simple motion along circular orbits assumes that the GC system being studied is spherically distributed, that the rotation axis lies in the plane of the sky and that the projected angular velocity is a function of radius only.

Looking carefully at Fig. 1 and Table 7, it is noticeable that even though the mean velocity of the GCs, which is -59 km s^{-1} , is similar to the systemic motion of NGC 6822, there is a preference for the GCs to have radial velocities more negative than the systemic velocity of their host. Thus, we retain v_{sys} as a free parameter instead of fixing in to the known systemic velocity of NGC 6822.

The velocity dispersion is assumed to possess a Gaussian form, and not to change as a function of projected radius. It is mathematically described as

$$\sigma^2 = (\Delta v)^2 + \sigma_0^2, \quad (9)$$

where σ_0 is the intrinsic velocity dispersion of the NGC 6822 GC system while Δv are the uncertainties in the radial velocity measurements.

Combining these two components, our kinematic model \mathcal{M} takes the following form,

$$p_{i,\mathcal{M}}(v_i, \Delta v_i | v_p, \sigma) = \frac{1}{\sqrt{2\pi}\sigma} e^{-\frac{(v_i - v_p)^2}{2\sigma^2}}, \quad (10)$$

where, as before, v_p is the systemic rotation described by equation (8), v_i are the observed heliocentric radial velocities of the GCs as presented in Table 7, and σ is a velocity dispersion as modelled by equation (9). Hence, the likelihood function for the above kinematic model is

$$p_{\mathcal{M}}(D|\Theta) = \mathcal{L}_{\mathcal{M}}(v, \Delta v, \theta | A, \theta_0, v_{\text{sys}}, \sigma_0) = \prod_i p_{i,\mathcal{M}}, \quad (11)$$

where v , Δv , and θ are the observed properties of the GCs, while A , θ_0 , v_{sys} and σ_0 are the model free parameters. We assume all priors for the free parameters to be flat.

Even though our model features four free parameters, we systematically calculate the likelihood function via a ‘brute-force’ method over a regular grid shaped by the parameter space of the kinematic model. The rotation amplitude A and the velocity dispersion parameters σ_0 range from 0 to 180 km s^{-1} and from 0 to 120 km s^{-1} , respectively, both with a resolution of 2 km s^{-1} . The axis of rotation, θ_0 , is searched from 0 to 2π rad with a step size of 0.15 rad, while the v_{sys} ranges from -200 to 45 km s^{-1} with a step size of 3 km s^{-1} .

With this set-up, however, it quickly became apparent that the algorithm could not properly converge with the position angle θ_0 allowed to vary as a free parameter. This is due to the unusual linear configuration of the GCs and the small size of the sample. Thus, in order to investigate the general kinematic properties of this GC system, we consider three constrained models. In the first scenario, we assume that if the GC system is rotating, the rotation axis is perpendicular to the GC alignment in the plane of the sky. This is a reasonable assumption given that rotating configurations usually

form flattened ellipsoids or discs. We refer to this model as to the ‘disc model’; in this scenario, the rotation axis would lie close to that inferred for the spheroid by Demers et al. (2006). In the second case, we fix θ_0 so that the rotation axis lies parallel to the apparent alignment of the GCs. This rotating ‘cigar model’ represents the case where the cluster system might share the rotation exhibited by the H I disc. Finally, we consider a scenario in which the GCs have kinematic properties separate from those of the spheroid and the H I disc, and exhibit no net rotation signal (i.e. v_p in equation 10 is set to zero and the problem is reduced to one with only two free parameters). This outcome would match that inferred by Hwang et al. (2014) from their kinematic study of SC1-SC4.

The results of our analysis are shown in Fig. 5. The general kinematic picture for this GC system is best described in the top-left panel, where the radial velocities of the GCs, corrected for the systemic motion of NGC 6822, are plotted against their position angles. In the same panel, the best-fitting rotation curves are overplotted, as described by equation (8), from our two scenarios which allow for rotation. The top-right panel shows the marginalized posterior probability distribution functions for the amplitude in these two rotation-enabled models. From these two panels, one can see that the disc model detects a low signature of net rotation, where the most likely amplitude is $12 \pm 10 \text{ km s}^{-1}$ with 68 per cent confidence. The probability of not having any net rotation signal is moderately high according to this model.

For the cigar model, on the other hand, a strong coherent rotation motion is inferred for the GC system. The most likely amplitude is $56 \pm 31 \text{ km s}^{-1}$ with 68 per cent confidence. This is quite a surprising result, since any rapidly rotating system ought naturally to form a flattened ellipsoid or disc structure. One possibility is that we are observing a dynamically young system; however, there is at present no evidence that this is true for the NGC 6822 GCs – for example one might expect to see a stellar stream or other tidal debris associated with the clusters. A small amount of experimentation revealed that this rotation signal is most likely a false detection driven by SC3 on the one hand, and SC6 and SC7 on the other. These objects have strong deviations from the mean systemic velocity in opposite directions, and happen to lie either side of the rotation axis when it is arranged parallel to the GC alignment. A small change in the assumed value of θ_0 in either direction strongly decreases the rotation signal. Because there are only a handful of GCs in the entire system, these few points are able to substantially influence the outcome of the calculation (and indeed we believe it was this numerical instability as θ_0 moves across the GC alignment that led our original four-parameter analysis to fail).

The lower-left and -right panels of Fig. 5 show the marginalized posterior distributions of the velocity dispersion θ_0 and the systemic velocity of NGC 6822 v_{sys} . Regardless of the assumed model, the most likely value of v_{sys} is always found to be between -59 and -60 km s^{-1} with 68 per cent confidence intervals of up to $\pm 13 \text{ km s}^{-1}$. This is in excellent agreement with the literature value for the systemic velocity of NGC 6822 from McConnachie (2012). The velocity dispersion is also well defined and does not show a large variation depending on the assumed model. Its most likely value is 22_{-16}^{+10} , 26_{-18}^{+12} , $27_{-19}^{+11} \text{ km s}^{-1}$ for the cigar model, disc model, and no-rotation model, respectively.

6 THE DYNAMICAL MASS OF NGC 6822

GCs have a long history of being used as probes for estimating the dynamical masses of the galaxies they surround (e.g. Federici, Marano & Fusi Pecci 1990; Federici et al. 1993; Evans & Wilkinson

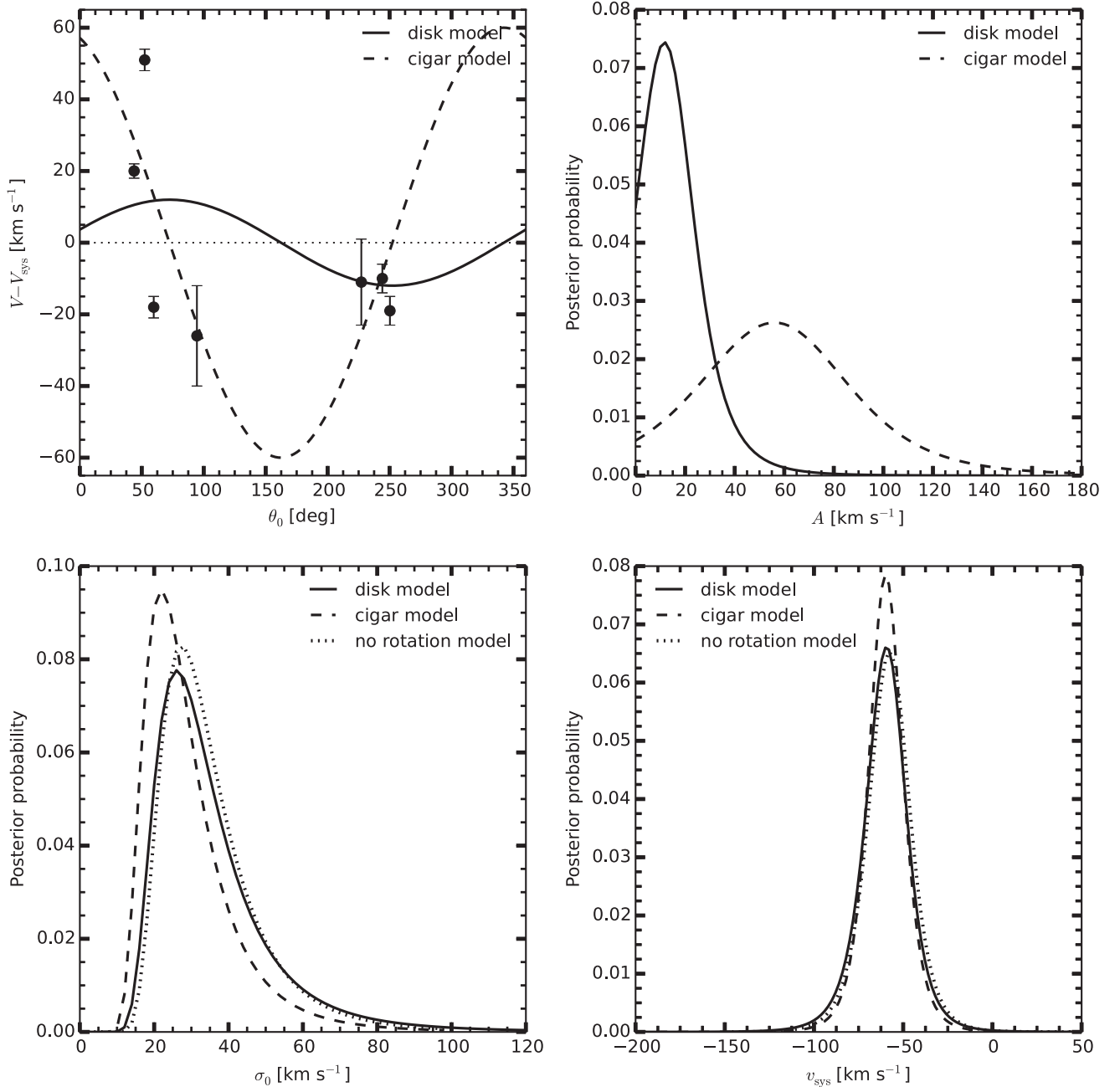


Figure 5. Top left: heliocentric radial velocity, corrected for the systemic motion of NGC 6822, versus position angle of the GCs on the sky. The most-likely rotation curves from our two rotation-enabled models are displayed. The disc model shows evidence for a low degree of net rotation, while the cigar model detects a more considerable rotation signal (although this is likely to be non-physical – see the text). Top right: the marginalized posterior distribution functions for rotation amplitude. The amplitude for the disc model peaks at $12 \pm 10 \text{ km s}^{-1}$, where the uncertainty corresponds to the 68 percent confidence interval. The amplitude for the cigar model peaks at $56 \pm 31 \text{ km s}^{-1}$. Lower left: the posterior functions for σ_0 , with peaks between $22\text{--}27 \text{ km s}^{-1}$ depending on the model used. In all three cases, the distributions are quite similar. Lower right: the posterior functions for v_{sys} . The functions have peaks between -59 and -60 km s^{-1} , which is in excellent agreement with the systemic velocity of NGC 6822 found in the literature (McConnachie 2012).

2000). We estimate the mass of NGC 6822 using its GCs as dynamical tracers following a common approach (e.g. Woodley et al. 2010; Veljanoski et al. 2013b, 2014). Since we considered three kinematic scenarios, we will use the results of each to estimate the mass of this dwarf galaxy. Solving the Jeans equations (Binney & Tremaine 1987) is one common approach. Since our rotation-enabled models found a varying degree of rotation in NGC 6822’s GC system, the total mass of this dwarf galaxy can be thought of as the sum

from two components: a rotating component, and a non-rotating pressure-supported component.

In the scenarios where a net rotation signature is detected, the rotation component, M_r , is simply calculated via

$$M_r = \frac{R_{\text{max}} A^2}{G}, \quad (12)$$

Table 8. Dynamical mass estimates for NGC 6822. See the text for details.

| Model | N_{GC} | α | $M_p (M_\odot)$ | $M_r (M_\odot)$ | $M_{total} (M_\odot)$ | $M/L(M_\odot/L_\odot)$ |
|-------------|----------|----------|--------------------------|-----------------------|--------------------------|------------------------|
| Disc | 6 | 0 | $4 \pm 2 \times 10^9$ | $3 \pm 6 \times 10^8$ | $4 \pm 2 \times 10^9$ | 40 |
| Disc | 6 | 0.55 | $3 \pm 1 \times 10^9$ | $3 \pm 6 \times 10^8$ | $4 \pm 1 \times 10^9$ | 40 |
| Cigar | 6 | 0 | $4 \pm 2 \times 10^{10}$ | $7 \pm 8 \times 10^9$ | $4 \pm 2 \times 10^{10}$ | 400 |
| Cigar | 6 | 0.55 | $3 \pm 1 \times 10^{10}$ | $7 \pm 8 \times 10^9$ | $4 \pm 2 \times 10^{10}$ | 400 |
| No rotation | 6 | 0 | $4 \pm 1 \times 10^9$ | | $4 \pm 1 \times 10^9$ | 40 |
| No rotation | 6 | 0.55 | $3 \pm 1 \times 10^9$ | | $3 \pm 1 \times 10^9$ | 30 |

where A is the rotation amplitude of the GC system, R_{\max} is the projected radius of the outermost constituent, and G is the gravitational constant.

The pressure supported mass component is determined using the ‘Tracer Mass Estimator’ (TME) devised by Evans et al. (2003). The TME is mathematically described as

$$M = \frac{C}{GN} \sum_{i=1}^N (v_i - v_{\text{sys}})^2 R_i, \quad (13)$$

where R is the projected radius from the centre of NGC 6822 for a given GC, v is the radial velocity of that GC, corrected for global rotation of the system, while v_{sys} is the systemic radial velocity of NGC 6822 itself. The index i loops over each GC in the sample of N GCs that have available radial velocities. The constant C is dependent on the shape of the underlying gravitational potential, α , assumed to be scale-free; the radial distribution of the mass tracers, assumed to be a power law with index γ ; and the anisotropy of the system. For an anisotropic system, it takes the form

$$C = \frac{4(\alpha + \gamma)}{\pi} \frac{4 - \alpha - \gamma}{3 - \gamma} \frac{1 - (r_{\text{in}}/r_{\text{out}})^{3-\gamma}}{1 - (r_{\text{in}}/r_{\text{out}})^{4-\alpha-\gamma}}. \quad (14)$$

For $\gamma = 3$, which is typical for a spheroidal stellar halo, C becomes

$$C = \frac{4(\alpha + 3)(1 - \alpha)}{\pi} \frac{\log(r_{\text{out}}/r_{\text{in}})}{1 - (r_{\text{in}}/r_{\text{out}})^{1-\alpha}}. \quad (15)$$

We consider cases where the slope of the underlying gravitational field, α , is set to 0, which assumes an isothermal halo potential for NGC 6822, and to 0.55, as for an NFW profile (Navarro, Frenk & White 1996; Watkins, Evans & An 2010). For r_{in} and r_{out} , we adopt the smallest and largest projected radii, respectively, exhibited by the GCs in our sample. As before, we include the velocity measured for SC4 by Mackey et al. (in preparation); however, we now exclude Hubble-VII from the analysis. If a tracer object having a projected radius of zero is used in the TME, it will produce a singularity in equation (15). Hubble-VII has a projected radius of just 0.13 kpc, which causes the constant C to increase anomalously.

Table 8 lists estimated masses for NGC 6822 within ~ 11 kpc, i.e. the projected radius of the most remote cluster (SC1). The uncertainties of the pressure-supported mass component are calculated via the jackknifing technique, while the uncertainties in the rotation component are only due to the error propagation of the detected amplitude in the relevant scenarios. Total masses estimated assuming the rotating disc and the non-rotating scenario yield similar results, which is unsurprising given the low degree of rotation detected in the disc model. For these cases, our measurements yield an estimated mass of $3\text{--}4 \times 10^9 M_\odot$, with a corresponding mass-to-light ratio of 30–40. These estimates are, however, smaller than those from Hwang et al. (2014) who found the mass of NGC 6822 within 11 kpc to be $7.5^{+4.5}_{-0.1} \times 10^9 M_\odot$ using four extended clusters only (SC1–SC4). The discrepancy in the estimates is likely due to

the large uncertainties in the velocity measurements, in combination with their smaller cluster sample. Still, our inferred mass-to-light ratio makes NGC 6822 a highly dark-matter-dominated system when compared to other dwarf galaxies of similar luminosity (e.g. Mateo 1998; McConnachie 2012; Kirby et al. 2014).

On the other hand, assuming the rotating cigar scenario we estimate the mass of NGC 6822 to be unreasonably high, about $4 \times 10^{10} M_\odot$, as indicated by the associated mass-to-light ratio. This adds further confidence that the high rotation signal found for this model is indeed a spurious detection.

It is worth acknowledging the caveats regarding the choice of α and γ parameters which feature in the TME. It is unclear whether their assumed values are indeed the right choices when estimating the mass of NGC 6822, thus adding additional uncertainty which is not formally included. Using the same method and identical assumptions, but only four extended GCs, Hwang et al. (2014) found the total mass of NGC 6822 enclosed within 11 kpc to be $7.5^{+4.5}_{-0.1} \times 10^9 M_\odot$. This most likely indicates that when dealing with a small samples such as this one, the inclusion or exclusion of a single data point can have an important effect on the results. For reference, using the kinematics of the H I gas, Weldrake et al. (2003) found a mass of $3.2 \times 10^9 M_\odot$ within ~ 5 kpc from the centre of NGC 6822.

7 DISCUSSION

GCs hosted by dwarf galaxies have been often found to have similar colours as GCs residing in the haloes of massive galaxies (e.g. Kundu & Whitmore 2001; Larsen, Forbes & Brodie 2001; Peng et al. 2006), and our results are in keeping with that. Comparing the optical colours of the GCs hosted by NGC 6822 with those located in the M31 outer halo, it is found that they are all mutually consistent, having $(V - I)_0 \sim 0.9$. This is also true for the GCs residing in the dwarf elliptical galaxies NGC 147 and NGC 185, located in the outer halo of M31 (Veljanoski et al. 2013a). In terms of their number, extended GCs dominate the population of NGC 6822 and there is a similarity between the mean absolute magnitudes of the extended GCs hosted by NGC 6822 and by M31, which are found to be $M_{V_0} = -5.9 \pm 0.7$ and $M_{V_0} = -5.6 \pm 0.7$, respectively. This may either be a fundamental property of extended GCs in general, or is a result of a bias arising from the faint limit of the survey imaging from which these objects were discovered. Given that some of them lie right at the detection limit, it is possible that the extended GCs have a larger range of absolute magnitudes than observed at present, but the currently available data allows only for the bright end of the luminosity function to be observed, resulting in the low spread and similar M_{V_0} values. None the less, the range of luminosities, structures, ages and metallicities spanned by the GCs in NGC 6822 is broadly consistent with that observed in the M31 halo population (e.g. Mackey et al. 2006, 2007, 2010a, 2013a,b; Alves-Brito et al. 2009; Sakari et al. 2015).

A number of studies have focused on GC systems hosted by dIrr galaxies outside the Local Group (e.g. Seth et al. 2004; Sharina, Puzia & Makarov 2005; Georgiev et al. 2006, 2008). They have found that GCs have typical optical colours within the range of $0.8 < (V - I)_0 < 1.1$, consistent with our measurements. In their study, Georgiev et al. (2008) searched for GCs in 19 nearby (2–8 Mpc) dIrr galaxies using archival *HST* data. The galaxies in their sample are members of dwarf galaxy associations only, without a dominant massive galaxy nearby. In the final sample of GC candidates that passed all their selection criteria, Georgiev et al. (2008) found an absence of objects having $(V - I)_0 < 1$ and $M_{V_0} \gtrsim -6$, i.e. faint and blue GCs. The authors claim that this effect is not due to a bias in their GC selection criteria, nor a bias due to the depth of the imaging, since they are able to identify clusters down to $M_{V_0} \simeq -4$. In NGC 6822 however, there are two extended GCs which satisfy these criteria. If such clusters were present in the sample of Georgiev et al. (2008), they were most likely missed because of their low surface brightness as well as their extended morphology.

We have considered three scenarios for the kinematic configuration of the GC system of NGC 6822. As indicated in the previous section, the so-called ‘disc’ and ‘no-rotation’ models are the most plausible of these. In the disc scenario our analysis provides tentative evidence for a weak rotation signal of amplitude $12 \pm 10 \text{ km s}^{-1}$, although clearly the case of zero rotation is not strongly excluded. It is interesting to note that this putative rotation would be in the same sense as that detected for stars in the underlying stellar spheroid by Demers et al. (2006), and also of a comparable magnitude (although we note that the Demers et al. measurements do not extend to particularly large galactocentric radii). Similarly, the velocity dispersion we infer for the GC system matches closely that determined from the Demers et al. (2006) kinematic sample, once the global rotation has been accounted for (see Hwang et al. 2014). These similarities suggest that the GC system of NGC 6822 is plausibly associated with its stellar spheroid, and indeed this would be consistent with the clear linear alignment of the GCs with the major axis of the spheroid.

Our third kinematic model – the so-called ‘cigar’ model – would imply a large systemic rotation amongst the GC population, although this is strongly sensitive to the precise orientation of the rotation axis. In this case, the sense of rotation would be opposite to that seen for the H I gas by de Blok & Walter (2000); furthermore, the inferred amplitude would be considerably larger than that seen for the H I gas at comparable distances along its major axis to those of the GCs. In the event this model was correct, it would appear that the GCs are not associated with the H I gas component of NGC 6822. The GCs could constitute a dynamically young system, in which case a search for a stellar stream along the major axis of the NGC 6822 spheroid may prove fruitful.

8 SUMMARY

In this paper, we present a uniform optical and near-IR photometric study of the GC system hosted by NGC 6822. The photometric measurements are used to estimate the ages and metallicities of the GCs via a colour–colour plot and an empirical colour–metallicity relation. All GCs are found to have ages consistent with 9 Gyr or older, in agreement with past studies. On the other hand, the clusters are found to exhibit a range of metallicities ($-1.6 \lesssim [\text{Fe}/\text{H}] \lesssim -0.4$), at odds with previous spectroscopic studies which found them to be very metal-poor ($[\text{Fe}/\text{H}] < -2.0$). Metallicity measurements from

resolved colour–magnitude diagrams would alleviate this tension, as would higher S/N spectroscopy.

We also modelled the kinematics of the GC system using new velocities for six members, two of which had no previous spectroscopic information. With the caveat at our results are based on a small sample, we found tentative evidence for a weak rotation signal of amplitude $12 \pm 10 \text{ km s}^{-1}$ in the case where the rotation axis sits perpendicular to the linear arrangement of the GCs on the sky. However, zero rotation is not strongly excluded by our analysis. If the rotation is real, then it would be in the same sense as inferred for the underlying stellar spheroid by Demers et al. (2006), and of a comparable amplitude. Furthermore, the velocity dispersion we determine for the GC system, $\sigma_0 = 26_{-12}^{+18} \text{ km s}^{-1}$ is very similar to that exhibited by stars in the spheroid (see Hwang et al. 2014). These similarities suggest to us that the GCs are plausibly associated with the spheroidal component of NGC 6822.

Finally, using the GCs as dynamical mass tracers, the total mass of NGC 6822 is recalculated to be roughly $(3\text{--}4) \times 10^9 M_\odot$, subject to various systematic effects due the small data sample and the implicit assumptions built into the TME. The corresponding mass-to-light ratio sits in the range $\sim 30\text{--}40$, implying that NGC 6822 hosts a substantial dark matter component.

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